

## New reaction rates for the destruction of ${}^7\text{Be}$ during big bang nucleosynthesis measured at CERN/n\_TOF and their implications on the cosmological lithium problem

A. Mengoni<sup>1,2</sup>, L.A. Damone<sup>3,4</sup>, M. Barbagallo<sup>5,3</sup>, O. Aberle<sup>5</sup>, V. Alcayne<sup>6</sup>, S. Amaducci<sup>7,8</sup>, J. Andrzejewski<sup>9</sup>, L. Audouin<sup>10</sup>, V. Babiano-Suarez<sup>11</sup>, M. Bacak<sup>5,12,13</sup>, S. Bennett<sup>14</sup>, E. Berthoumieux<sup>13</sup>, D. Bosnar<sup>15</sup>, A.S. Brown<sup>16</sup>, M. Busso<sup>17,18</sup>, M. Caamaño<sup>19</sup>, L. Caballero<sup>11</sup>, M. Calviani<sup>5</sup>, F. Calviño<sup>20</sup>, D. Cano-Ott<sup>6</sup>, A. Casanovas<sup>20</sup>, F. Cerutti<sup>5</sup>, E. Chiaveri<sup>14,21,5</sup>, N. Colonna<sup>3</sup>, G.P. Cortés<sup>20</sup>, M.A. Cortés-Giraldo<sup>21</sup>, L. Cosentino<sup>7</sup>, S. Cristallo<sup>17,22</sup>, P.J. Davies<sup>14</sup>, M. Diakaki<sup>23</sup>, M. Dietz<sup>24</sup>, C. Domingo-Pardo<sup>11</sup>, R. Dressler<sup>25</sup>, Q. Ducasse<sup>26</sup>, E. Dupont<sup>13</sup>, I. Durán<sup>19</sup>, Z. Eleme<sup>27</sup>, B. Fernández-Domínguez<sup>19</sup>, A. Ferrari<sup>5</sup>, I. Ferro-Gonçalves<sup>28</sup>, P. Finocchiaro<sup>7</sup>, V. Furman<sup>29</sup>, R. Garg<sup>24</sup>, A. Gawlik<sup>9</sup>, S. Gilardoni<sup>5</sup>, K. Göbel<sup>30</sup>, E. González-Romero<sup>6</sup>, C. Guerrero<sup>21</sup>, F. Gunsing<sup>13</sup>, S. Heinitz<sup>25</sup>, J. Heyse<sup>31</sup>, D.G. Jenkins<sup>16</sup>, E. Jericha<sup>12</sup>, U. Jiri<sup>25</sup>, A. Junghans<sup>32</sup>, Y. Kadi<sup>5</sup>, F. Käppeler<sup>33</sup>, A. Kimura<sup>34</sup>, I. Knapová<sup>35</sup>, M. Kokkoris<sup>23</sup>, Y. Kopatch<sup>29</sup>, M. Krtička<sup>35</sup>, D. Kurtulgil<sup>30</sup>, I. Ladarescu<sup>11</sup>, C. Lederer-Woods<sup>24</sup>, J. Lerendegui-Marco<sup>21</sup>, S.-J. Lonsdale<sup>24</sup>, D. Macina<sup>5</sup>, A. Manna<sup>2,36</sup>, T. Martínez<sup>6</sup>, A. Masi<sup>5</sup>, C. Massimi<sup>2,36</sup>, P.F. Mastinu<sup>37</sup>, M. Mastroianni<sup>5,14</sup>, E. Mauger<sup>25</sup>, A. Mazzone<sup>3,38</sup>, E. Mendoza<sup>6</sup>, V. Michalopoulou<sup>5,23</sup>, P.M. Milazzo<sup>39</sup>, M.A. Millán-Callado<sup>21</sup>, F. Mingrone<sup>5</sup>, J. Moreno-Soto<sup>13</sup>, A. Musumarra<sup>7,8</sup>, A. Negret<sup>40</sup>, F. Ogállar<sup>41</sup>, A. Oprea<sup>40</sup>, N. Patronis<sup>27</sup>, A. Pavlik<sup>42</sup>, J. Perkowski<sup>9</sup>, C. Petrone<sup>40</sup>, L. Piersanti<sup>17,22</sup>, E. Pirovano<sup>26</sup>, I. Porras<sup>41</sup>, J. Praena<sup>41</sup>, J.M. Quesada<sup>21</sup>, D. Ramos Doval<sup>10</sup>, R. Reifarth<sup>30</sup>, D. Rochman<sup>25</sup>, C. Rubbia<sup>5</sup>, M. Sabaté-Gilarte<sup>21,5</sup>, A. Saxena<sup>43</sup>, P. Schillebeeckx<sup>31</sup>, D. Schumann<sup>25</sup>, A. Sekhar<sup>14</sup>, A.G. Smith<sup>14</sup>, N. Sosnin<sup>14</sup>, P. Sprung<sup>25</sup>, A. Stamatopoulos<sup>23</sup>, G. Tagliente<sup>3</sup>, J.L. Tain<sup>11</sup>, A.E. Tarifeño-Saldivia<sup>20</sup>, L. Tassan-Got<sup>5,23,10</sup>, B. Thomas<sup>30</sup>, P. Torres-Sánchez<sup>41</sup>, A. Tsinganis<sup>5</sup>, S. Urluss<sup>5,32</sup>, S. Valenta<sup>35</sup>, G. Vannini<sup>2,36</sup>, V. Variale<sup>3</sup>, P. Vaz<sup>28</sup>, A. Ventura<sup>2</sup>, D. Vescovi<sup>17,44</sup>, V. Vlachoudis<sup>5</sup>, R. Vlastou<sup>23</sup>, A. Wallner<sup>45</sup>, P.J. Woods<sup>24</sup>, T.J. Wright<sup>14</sup>, and P. Žugec<sup>15</sup>

<sup>1</sup>Agenzia nazionale per le nuove tecnologie, l'energia e lo sviluppo economico

<sup>2</sup>Istituto Nazionale di Fisica Nucleare, Sezione di Bologna, Italy

<sup>3</sup>Istituto Nazionale di Fisica Nucleare, Bari, Italy

<sup>4</sup>Dipartimento di Fisica, Università degli Studi di Bari, Italy

<sup>5</sup>European Organization for Nuclear Research (CERN), Switzerland

<sup>6</sup>Centro de Investigaciones Energéticas Medioambientales y Tecnológicas (CIEMAT), Spain

<sup>7</sup>INFN Laboratori Nazionali del Sud, Catania, Italy

<sup>8</sup>Dipartimento di Fisica e Astronomia, Università di Catania, Italy

<sup>9</sup>University of Lodz, Poland

<sup>10</sup>IPN, CNRS-IN2P3, Univ. Paris-Sud, Université Paris-Saclay, Orsay, France

<sup>11</sup>Instituto de Física Corpuscular, CSIC - Universidad de Valencia, Spain

<sup>12</sup>Technische Universität Wien, Austria

<sup>13</sup>CEA Saclay, Irfu, Université Paris-Saclay, Gif-sur-Yvette, France

<sup>14</sup>University of Manchester, United Kingdom

<sup>15</sup>Department of Physics, Faculty of Science, University of Zagreb, Croatia

<sup>16</sup>University of York, United Kingdom

<sup>17</sup>Istituto Nazionale di Fisica Nucleare, Perugia, Italy

<sup>18</sup>Dipartimento di Fisica e Geologia, Università di Perugia, Italy

<sup>19</sup>University of Santiago de Compostela, Spain

<sup>20</sup>Universitat Politècnica de Catalunya, Spain

<sup>21</sup>Universidad de Sevilla, Spain

<sup>22</sup>Istituto Nazionale di Astrofisica - Osservatorio Astronomico d'Abruzzo, Italy

<sup>23</sup>National Technical University of Athens, Greece

<sup>24</sup>School of Physics and Astronomy, University of Edinburgh, United Kingdom

<sup>25</sup>Paul Scherrer Institut (PSI), Villigen, Switzerland

<sup>26</sup>Physikalisch-Technische Bundesanstalt (PTB), Braunschweig, Germany

<sup>27</sup>University of Ioannina, Greece

<sup>28</sup>Instituto Superior Técnico, Lisbon, Portugal

<sup>29</sup>Joint Institute for Nuclear Research (JINR), Dubna, Russia

<sup>30</sup>Goethe University Frankfurt, Germany

<sup>31</sup>European Commission, Joint Research Centre, Geel, Belgium

<sup>32</sup>Helmholtz-Zentrum Dresden-Rossendorf, Germany

<sup>33</sup>Karlsruhe Institute of Technology, Karlsruhe, Germany

<sup>34</sup>Japan Atomic Energy Agency (JAEA), Tokai-mura, Japan

<sup>35</sup>Charles University, Prague, Czech Republic

<sup>36</sup>Dipartimento di Fisica e Astronomia, Università di Bologna, Italy

<sup>37</sup>Istituto Nazionale di Fisica Nucleare, Sezione di Legnaro, Italy

<sup>38</sup>Consiglio Nazionale delle Ricerche, Bari, Italy

<sup>39</sup>Istituto Nazionale di Fisica Nucleare, Trieste, Italy

<sup>40</sup>Horia Hulubei National Institute of Physics and Nuclear Engineering (IFIN-HH), Bucharest

<sup>41</sup>University of Granada, Spain

<sup>42</sup>University of Vienna, Faculty of Physics, Vienna, Austria

<sup>43</sup>Bhabha Atomic Research Centre (BARC), India

<sup>44</sup>Gran Sasso Science Institute (GSSI), L'Aquila, Italy

<sup>45</sup>Australian National University, Canberra, Australia

**Abstract.** New measurements of the  ${}^7\text{Be}(n,\alpha){}^4\text{He}$  and  ${}^7\text{Be}(n,p){}^7\text{Li}$  reaction cross sections from thermal to keV neutron energies have been recently performed at CERN/n\_TOF. Based on the new experimental results, astrophysical reaction rates have been derived for both reactions, including a proper evaluation of their uncertainties in the thermal energy range of interest for big bang nucleosynthesis studies. The new estimate of the  ${}^7\text{Be}$  destruction rate, based on these new results, yields a decrease of the predicted cosmological  ${}^7\text{Li}$  abundance insufficient to provide a viable solution to the cosmological lithium problem.

## 1 Introduction

A few neutron-induced reactions are important in the processes leading to the formation of the first elements at the very beginning of our universe, during the so-called big bang nucleosynthesis (BBN) era, spanning from a few seconds to a few minutes time duration and thermal energies from  $\sim 100$  keV down to a few keV. Amongst these, the (n,p) and (n, $\alpha$ ) reactions on  ${}^7\text{Be}$  play a key role, in particular for the determination of the abundance of primordial lithium. Considering that over 95% of the lithium resulting from the BBN is the product of the electron-capture decay of  ${}^7\text{Be}$ , the production and destruction mechanisms of this isotope are key elements in the determination of the primordial  ${}^7\text{Li}$  abundance, which is over-produced by BBN models by a factor 2-3 (the cosmological lithium problem, CLiP). While the  ${}^7\text{Be}$  production mechanisms, mostly going through the  ${}^3\text{H}(\alpha,\gamma){}^7\text{Be}$  reaction, have been thoroughly studied, the destruction channels have received relatively less attention and the related reaction rates have been based on old measurements, often complemented by theoretical assumptions. To verify the possibility of solution of the CLiP and to improve the confidence on the predictions of the BBN lithium yield, the n\_TOF Collaboration have recently performed measurements of the  ${}^7\text{Be}(n,\alpha){}^4\text{He}$  and  ${}^7\text{Be}(n,p){}^7\text{Li}$  reaction cross sections [1, 2], the main reaction mechanisms, leading to the the destruction of  ${}^7\text{Be}$  at BBN temperatures.

## 2 Experiments

Both cross section measurements were performed at the second experimental area of the n\_TOF facility at CERN [3]. High purity material was produced at the Paul Scherrer Institute (PSI), extracting 200 GBq of  ${}^7\text{Be}$  from the water cooling system of the SINQ spallation source [5].

For the  ${}^7\text{Be}(n,\alpha){}^4\text{He}$  measurement, two samples with  $\approx 18$  GBq of activity each ( $1.4 \mu\text{g}$  of  ${}^7\text{Be}$ ) were produced.

They were sandwiched with  $3 \times 3 \text{ cm}^2$  active area and  $140 \mu\text{m}$  thickness silicon detectors and inserted directly into the n\_TOF neutron beam for irradiation. Strong rejection of background events was possible because of the time-of-flight technique coupled to the low duty-cycle of the primary beam of the n\_TOF facility. Coincidence signals for protons from the (n,p) channel,  $\gamma$ -rays from  ${}^7\text{Be}$  activity and  $\alpha$ 's from the  $n+{}^7\text{Li} \rightarrow {}^8\text{Li} (\beta^-, 840 \text{ ms}) \rightarrow {}^8\text{Be}^* \rightarrow 2\alpha$  reaction were excluded in the data analysis.

For the  ${}^7\text{Be}(n,p){}^7\text{Li}$  experiment, the  ${}^7\text{Be}$  material has been implanted on suited backing at CERN/ISOLDE-GPS separator and RILIS facilities using a 30 keV ( $\approx 45 \text{ nA}$ )  ${}^7\text{Be}$  beam. A silicon telescope, with 20 and 300 mm,  $5 \times 5 \text{ cm}^2$  strip devices for  $\Delta E$  and E detection respectively, was used in the measurement [6]. The procedure adopted demonstrated for the first time the feasibility of neutron measurements on samples produced at radioactive ion beam facilities.

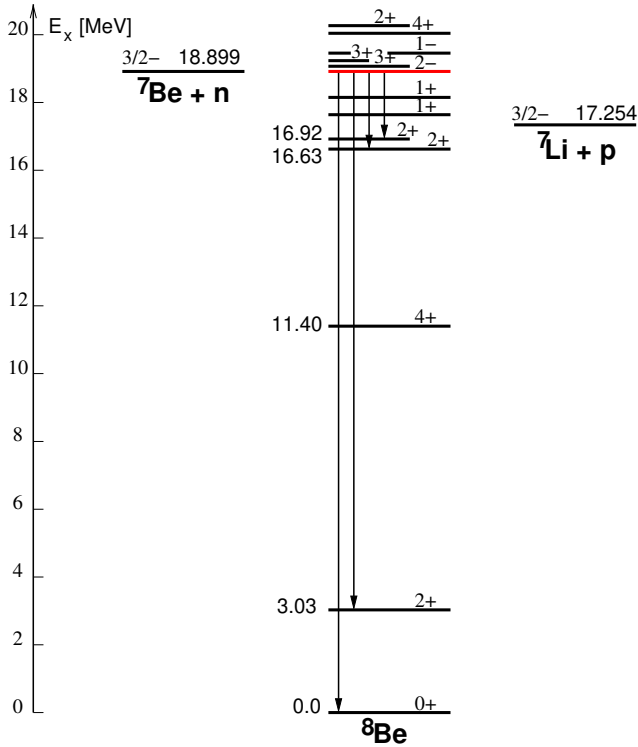
## 3 Results and implications

All the results of the measurement are reported in the references [1, 2]. Model interpretation, evaluation procedures and numerical tables (including uncertainties) of the measured cross sections are available online on the n\_TOF Collaboration twiki website [4]. The published data of both measurements are already available in the EXFOR database as well.

### 3.1 ${}^7\text{Be}(n,\alpha){}^4\text{He}$

The reaction process, induced by low-energy s-wave neutrons, is dominated by the  $2^-$  state located only a few keV above the neutron separation energy in  ${}^8\text{Be}$ , at  $E_x \approx 19 \text{ MeV}$  (see Figure 1). A direct  $2\alpha$ -breakup of this state is not allowed and, at these excitation energies, the reaction mechanism is dominated by the (n, $\gamma\alpha$ ) process.

The cross section for the  $\alpha$ 's emitted from the doublet  $2^+$  states at  $\approx 16.8 \text{ MeV}$  in  ${}^8\text{Be}$ , following the capture  $\gamma$ -ray transitions, was derived from the measurement.



**Figure 1.** Energy levels of  $^8\text{Be}$  in the energy range of interest for the present work.

The  $1/v$  behaviour of the cross section can be interpreted as a direct radiative capture process as well as as a compound resonance reaction mechanisms. For the first case, a model prediction can be made for all the allowed  $(n,\gamma\alpha)$  E1 transitions and, therefore, the total  $(n,\alpha)$  cross section can be derived. The resulting total  $(n,\alpha)$  cross section, complemented with data from time-reversal and other reaction channels in the higher energy region above  $E_n \approx 50$  keV [4], can be integrated over the energy range of interest for BBN network calculations in a proper temperature grid. The results can be represented accurately by the following expression of the reaction rate

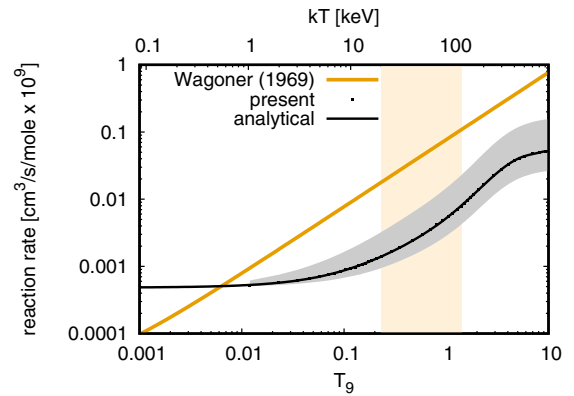
$$N_A \langle \sigma v \rangle = a_0(1 + a_1 T_9^{1/2} + a_2 T_9 + a_3 T_9^{3/2} + a_4 T_9^2 + a_5 T_9^{5/2} + a_6 T_9^3 + a_7 T_9^{7/2} + a_8 T_9^4 + a_9 T_9^{9/2} + a_{10} T_9^5) \quad (1)$$

in units of  $\text{cm}^3/\text{s}/\text{mole}$  when  $a_0 = 4.810 \times 10^5$ ,  $a_1 = -0.226$ ,  $a_2 = 5.301$ ,  $a_3 = 11.249$ ,  $a_4 = -18.940$ ,  $a_5 = 13.539$ ,  $a_6 = -0.133$ ,  $a_7 = -0.591$ ,  $a_8 = -1.144$ ,  $a_9 = 0.731$  and  $a_{10} = -0.094^1$ .

### 3.2 $^7\text{Be}(n,p)^7\text{Li}$

The measured cross section turned out to be higher than previously known, in particular at low neutron energies, up to  $\approx 35$  keV. The  $^7\text{Be}(n,p)^7\text{Li}$  measured cross section, complemented with data from the time-reversal channel  $^7\text{Li}(p,n)^7\text{Be}$ , has been fitted using single-level Breit-Wigner formalism with nine states above the neutron separation energy of  $^8\text{Be}$ , in order to fully cover the energy range of interest for BBN calculations. The resulting cross

<sup>1</sup>with respect to the rate published in [1], this expression includes additional terms in the expansion, making it valid up to  $T_9 = 10$ .

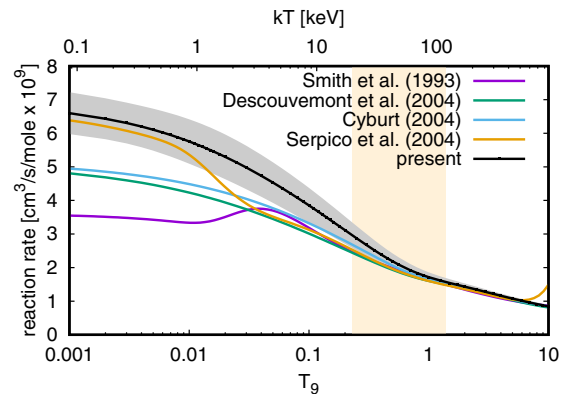


**Figure 2.**  $^7\text{Be}(n,\alpha)^4\text{He}$  rate is shown in comparison with the previously adopted rate of Wagoner [7]. The uncertainty associated with the presently determined rate is shown by the corresponding grey band. The temperature range of interest for the BBN is indicated by the vertical band

section has been integrated over the entire energy range to produce a reaction rate valid in the proper temperature range of interest for BBN network calculations

$$N_A \langle \sigma v \rangle = a_0(1 + a_1 T_9^{1/2} + a_2 T_9 + a_3 T_9^{3/2} + a_4 T_9^2 + a_5 T_9^{5/2}) + a_6 \left( \frac{1}{1 + 13.076 T_9} \right)^{3/2} + a_7 T_9^{-3/2} e^{-b_0/T_9} \quad (2)$$

in units of  $\text{cm}^3/\text{s}/\text{mole}$  when  $a_0 = 6.805 \times 10^9$ ,  $a_1 = -1.971$ ,  $a_2 = 2.042$ ,  $a_3 = -1.069$ ,  $a_4 = 0.271$ ,  $a_5 = -0.027$ ,  $a_6 = 1.961 \times 10^8$ ,  $a_7 = 2.890 \times 10^7$  and  $b_0 = 0.281$ .



**Figure 3.** Comparison of the reaction rates for the  $^7\text{Be}(n,p)^7\text{Li}$  reaction of the present work with some of the commonly adopted rates ([8–11]). The uncertainty associated with the presently determined rate is shown by the corresponding grey band. The temperature range of interest for BBN is indicated by the vertical band.

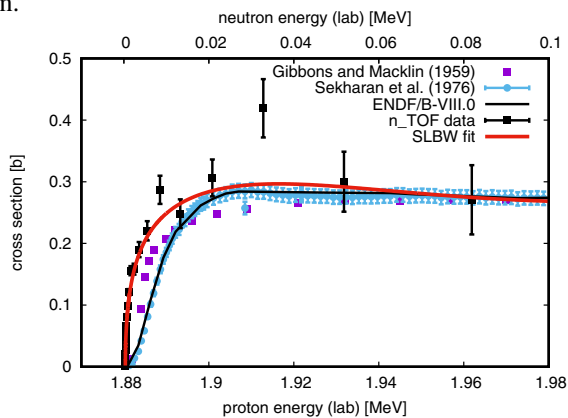
The new estimate of the  $^7\text{Be}$  destruction rates, based on the new  $n_{\text{TOF}}$  experimental results, can be used in BBN network calculations to estimate their impact on the lithium yield. Details on these calculations are provided in the references [1, 2, 4]. The BBN calculations have been performed adopting a neutron average life-time of

**Table 1.** Results of the BBN network calculation for the relevant main observables. Present rates refers only to the two rates evaluated in the present work. All the other network rates are adopted as described in [4].

	$Y_p$	D/H [ $10^{-5}$ ]	${}^3\text{He}/\text{H}$ [ $10^{-5}$ ]	${}^7\text{Li}/\text{H}$ [ $10^{-10}$ ]
with standard rates	0.246	2.43	1.08	5.46
using present rates ( $\eta_{10} = 6.09$ )	0.246	2.43	1.08	5.26
using present rates ( $5.8 \leq \eta_{10} \leq 6.6$ )	0.246	2.43	1.08	4.73 - 6.23
observations [15]	$0.245 \pm 0.003$	$2.569 \pm 0.027$	-	$1.6 \pm 0.3$

$\tau_n = 880.2$  s and  $N_\nu = 3$  neutrino species. The baryon-to-photon number density ratio in units of  $10^{-10}$ ,  $\eta_{10}$ , has been allowed to vary within the range established by the concordance of observation of primordial  ${}^4\text{He}$  and deuterium as evaluated in the review of the most recent Particle Data Group publication [15]. The results of the BBN calculation for the main observables are shown in the Table 1.

A decrease of the predicted cosmological lithium abundance (relative to H), from 5.46 to 5.26 in units of  $10^{-10}$  is predicted when using the new rates shown above. This is insufficient to provide a viable solution to the CLiP, leaving all alternative physics and astronomical scenarios open.



**Figure 4.**  ${}^7\text{Li}(p,n){}^7\text{Be}$  cross section, near the 1.88 MeV threshold, derived from time reversal invariance applied to the  $n\_TOF$   ${}^7\text{Be}(n,p){}^7\text{Li}$  data. Comparison is shown to the (p,n) measurements of Gibbons and Macklin [12] and Sekharan *et al.* [13] and to evaluated data file ENDF/B-VIII.0 [14]. The single-level Breit-Wigner cross section, calculated with the resonance parameters provided in [4], is plotted as well.

An additional outcome of the measurements performed at  $n\_TOF$  can be pointed out. The  ${}^7\text{Be}(n,p){}^7\text{Li}$  cross section data can be used to reconstruct the time reversal cross section of the  ${}^7\text{Li}(p,n){}^7\text{Be}$  reaction. In fact, the (n,p) data can provide the best cross section near the 1.88

MeV threshold, because the (n,p) channel has no threshold and the cross section has been measured in our experiment for neutron energies as low as meV. The results are shown in figure 4. In spite of the limited counting rates that causes fluctuations for neutron energies above 20 keV, this result is particularly relevant for all applications of the  ${}^7\text{Li}(p,n){}^7\text{Be}$  reaction as neutron source.

## References

- [1] M. Barbagallo *et al.* (The  $n\_TOF$  Collaboration), *Phys. Rev. Lett.* **117**, 152701, 2016
- [2] L.A. Damone *et al.* (The  $n\_TOF$  Collaboration), *Phys. Rev. Lett.* **121**, 042701, 2018
- [3] C. Weiss *et al.* (The  $n\_TOF$  Collaboration), *Nucl. Instr. Meth. Phys. Res. A* **799**, 2015, 90
- [4] The twiki public pages of the  $n\_TOF$  Collaboration: <http://twiki.cern.ch/NTOFPublic>.
- [5] E.A. Maugeri *et al.* (The  $n\_TOF$  Collaboration), *Nucl. Instr. Meth. Phys. Res. A* **889**, 138, 2018
- [6] M. Barbagallo *et al.* (The  $n\_TOF$  Collaboration), *Nucl. Instr. Meth. Phys. Res. A* **887**, 27, 2018
- [7] R.V. Wagoner, *ApJS*, **18**, 247, 1969
- [8] M.S. Smith, L. Kawano, and L.H. Malaney, *ApJ* **85**, 219, 1993
- [9] P. Descouvemont *et al.*, *Atomic Data and Nuclear Data Tables* **88**, 203, 2004
- [10] R.H. Cyburt, *Phys. Rev. D* **70**, 023505, 2004
- [11] P.D. Serpico *et al.*, *Journal of Cosmology and Astroparticle Physics* **12**, 10, 2004
- [12] J.H. Gibbons and R.L. Macklin, *Phys. Rev.* **114**, 571, 1959
- [13] K.K. Sekharan *et al.*, *Nucl. Instr. Meth. Phys. Res.* **133**, 253, 1976
- [14] A. Hermanne *et al.*, *Nuclear Data Sheets* **148**, 338, 2018
- [15] C. Patrignani *et al.* (Particle Data Group), *Chin. Phys. C* **40**, 100001, 2016